

STELLAR MOTIONS ALONG $l = 60 - 240$ AND $l = 150 - 330$ INHOMOGENEOUS FREQUENCY DISTRIBUTION OF SPECTRAL TYPES

S. MENGE DE FREITAS
Observatório do Valongo
Departamento de Astronomia - UFRJ

Abstract - Applications of the χ^2 test to the frequency distribution of spectral types show a slight predominance of A, F and G stars running approximately along $l = 60 - 240$; along the orthogonal direction, there seems to occur a preponderance of K and M stars.

Resumo - Aplicando-se o teste de Qui Quadrado às distribuições de frequências dos tipos espectrais, verifica-se uma certa predominância de estrelas A, F e G movendo-se aproximadamente ao longo de $l = 60 - 240$; percorrendo a direção ortogonal a esta, parece ocorrer uma preponderância de estrelas K e M.

INTRODUCTION

Usually studies on stellar kinematics seek to describe behaviours followed by groups (spectral types, luminosity classes, MK classifications) of stars, as well as to evaluate characteristic parameters of their motions. Therefore, the customary procedure is to take the stars themselves as starting-points. However, the present paper takes, not the stars, but the paths followed by them, as starting-points.

In our three previous papers on velocities of nearby stars (Freitas 1980, 1982, 1985) we have used components denoted by L (positive in the direction of Galactic Longitude $l=60$) and T (positive in $l=330$). In the first study, a sample of 726 stars, without sharp localization boundaries, was used. In the second and third papers, the sample was confined to a square centered in the Sun, with 36 pc on each side; this reduced the sample to 631 stars. Distances to the galactic plane were not considered in the three papers. Referring again to their first one, we have considered stars for which $|T| > 3|L|$, $|T| > 4|L|$ and $|T| > 5|L|$, in order to see how

perpendicular crossings of the longitude line $60 - 240$, in both senses, could affect dispersions in the T component. The imposition of such conditions gave us subsamples with, respectively, 155, 111 and 93 stars.

In this study, we consider not only the three above conditions (confinement to the squares reducing subsamples to 137, 101 and 83 stars respectively) but also stars for which $|L| > 3|T|$, $|L| > 4|T|$ and $|L| > 5|T|$ (109, 79 and 59 stars). Factors 3, 4 and 5 imply in path deviations to the largest component no greater than 18.43, 14.04 and 11.31, respectively (tables 1 through 6 only mention the first decimal).

Therefore, we consider two groups of paths, one approximately along longitude 150 or 330, and the other approximately along 60 or 240. Of course the paths here considered are not galactic orbits themselves, but only very tiny section of them.

THE PROPOSED PROBLEM

In papers I and II, the T components (approximately perpendicular to the direction

of the local spiral arm) have shown and oddity in their behaviours, absolutely not shared by the L components. This result suggests, at least, an inspection on stars running approximately along the directions which define L and T.

Even though our stellar samples occupy a very small section of the Galaxy, we must recall that, as Woolley et al. (1969) observe, "these stars, although close to us at this moment, in many cases describe orbits in the Galaxy which take them to distances of one or more kiloparsecs from the Sun. Gliese's Catalogue is thus a sample of a much larger volume of the Galaxy than the space now occupied". We have then, in our vicinity, stars which are called "visitors" by Trumpler and Weaver (1953).

We should expect that such visitors must include a very large fraction of old stars. Their galactic orbits may have rosette shapes (Mihalas, 1968), and very hardly young stars could have suffered enough collisions as to deform in such way their original orbits. Therefore, we should expect to find a larger fraction of old stars running along our T component than running along L.

The sequence of spectral types from A through M, even in Luminosity Class V (the vast majority in our samples), is a sequence of decreasing mass. Thus, it may be considered as a sequence of increasing longevity. Therefore, following this sequence, we increase the probability of finding old stars among the spectral types. Considering the indicators of old age and the data used in Table II of paper III, we see that, in a wider sample, the fraction of old stars grows from G to M spectral types.

We could expect, then, to find more K and M stars running along the T component than along L. A simple counting in our samples show that, in fact, such predominance occurs.

On the other hand, from a sample of 137 stars which run along the T component itself, 150 or 330 \pm 18.4, we have 116 stars of spectral types G, K and M with data for their orbital eccentricities (Woolley et al., 1969). The averages for this parameter and respective standard deviations are

2 G stars: $e = 0.11$ $s = 0.05$

46 K stars $e = 0.11$ $s = 0.07$

48 M stars $e = 0.11$ $s = 0.10$.

Both averages and dispersions suggest an increase on the number of old stars from G through M, running along the above direction.

Considering the above items, a statistical investigation seems to be justifiable, even for the small samples we have. Therefore, taking the paths 150 or 330 and 60 or 240 (with the tolerance deviations already explained) as starting-points, we try to verify if the spectral distribution of stars which run along these directions is homogeneous or not.

THE METHOD

We have used the χ^2 test in order to see whether the paths and spectral types are independent or not. Tables 1 to 6 refer to the cases for which $\alpha < 0.100$ without applying Yates' Continuity Correction (YCC); for Tables 2 through 6 the significance level α obtained with the correction is also given (unchanged for Table 2).

In Table 1 we have one expected frequency smaller than 5; however, this Table has 10 cells, and one case of expected frequency with value between 2 and 5, even among 5 cells, does not impair the reliability of the χ^2 test (Choi, 1968, p. 109).

Different groupings of spectral types and one case of exclusion of G stars (to give a little more weight to extremes) are presented, for confirmation and comparison.

CONCLUSIONS

Among stars in the solar neighborhood which run approximately along $l = 60 - 240$, there seems to occur a predominance of A, F and G spectral types, while the direction 150-330 presents a preponderance of K and M types. (This is to be considered only as an indicative result due to smallness of the samples here considered). Future work on larger samples may or not confirm this dichotomy.

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CONTINGENCY TABLES FOR DIRECTIONS AND SPECTRAL TYPES

Table 1

DIR S.P	60° or 24°	± 18.4°	150° or 330°	± 18.4°	TOTALS
A	8	(4.43)	2	(2.57)	10
f	13	(11.96)	14	(15.04)	27
G	27	(22.60)	24	(28.40)	51
K	30	(43.56)	48	(43.44)	78
M	31	(35.45)	49	(44.55)	80
Totals	109			137	246
$\alpha = 0.100$					

Table 2

DIR S.P	60° or 240°	± 18.4°	150° or 330°	± 18.4°	TOTALS
A + F + G	48	(38.99)	40	(49.01)	88
K + M	61	(70.01)	97	(87.99)	158
Totals	109			137	246
$\alpha = 0.025$					
$\alpha = 0.025$ with YCC					

Table 3

DIR S.P	60° or 240°	± 18.4°	150° or 330°	± 18.4°	TOTALS
A + F	21	(16.39)	16	(20.61)	37
G + K + M	88	(92.61)	121	(116.39)	209
Totals	109			137	246
$\alpha = 0.100$					
$\alpha = 0.100$ with YCC					

Table 4

DIR S.P	60° or 240°	± 18.4°	150° or 330°	± 18.4°	TOTALS
A+F	21	(15.56)	16	(21.44)	37
K+M	61	(66.44)	97	(91.56)	158
Totals	82			113	195
$\alpha = 0.50$ $\alpha = 0.100$ with YCC					

Table 5

DIR S.P	60° or 240°	± 14°	150° or 330°	± 14°	TOTALS
A+F+G	31	(24.58)	25	(31.42)	56
K+M	48	(54.42)	76	(69.58)	124
Totals	79			101	180
$\alpha = 0.50$ $\alpha = 0.100$ with YCC					

Table 6

DIR S.P	60° or 240°	± 11.3°	150° or 330°	± 11.3°	TOTALS
A+F+G	24	(19.11)	22	(26.89)	46
K+M	35	(39.89)	61	(56.11)	96
Totals	59			83	142
$\alpha = 0.10$ $\alpha = 0.100$ with YCC					